



Polarized Kink Waves in Magnetic Elements: Evidence for Chromospheric Helical Waves

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Abstract

In recent years, new high spatial resolution observations of the Sun's atmosphere have revealed the presence of a plethora of small-scale magnetic elements down to the resolution limit of the current cohort of solar telescopes (~ 100 – 120 km on the solar photosphere). These small magnetic field concentrations, due to the granular buffeting, can support and guide several magnetohydrodynamic wave modes that would eventually contribute to the energy budget of the upper layers of the atmosphere. In this work, exploiting the high spatial and temporal resolution chromospheric data acquired with the Swedish 1 m Solar Telescope, and applying the empirical mode decomposition technique to the tracking of the solar magnetic features, we analyze the perturbations of the horizontal velocity vector of a set of chromospheric magnetic elements. We find observational evidence that suggests a phase relation between the two components of the velocity vector itself, resulting in its helical motion.

Key words: magnetohydrodynamics (MHD) – Sun: chromosphere – Sun: oscillations

1. Introduction

Small-scale magnetic elements (SSMEs) with diameters of the order of a few hundred km are ubiquitous in the lower solar atmosphere (Lagg et al. 2010; Bonet et al. 2012; Morton et al. 2012; Stangalini 2014). Interestingly, they play a significant role in the energy budget of the chromosphere, acting as magnetic conduits for magnetohydrodynamic waves (De Pontieu et al. 2004; Jefferies et al. 2006). Indeed, under the forcing action of the photospheric convection, SSMEs are continuously pushed, pulled, advected, and dispersed over the solar surface (see, for instance, Berger et al. 1998; Keys et al. 2011; Chitta et al. 2012; Giannattasio et al. 2013, 2014, and references therein). At the same time, different magnetohydrodynamic (MHD) wave modes propagating along these waveguides are also excited (e.g., magnetoacoustic, kink, and sausage, Alfvén; Roberts & Webb 1978; Spruit 1981; Edwin & Roberts 1983; Roberts 1983; Musielak et al. 1989; Steiner et al. 1998; Hasan et al. 2003; Musielak & Ulmschneider 2003; Fedun et al. 2011; Nutto et al. 2012; Vigeesh et al. 2012). In this regard, many authors have reported the presence of a plethora of waves in SSMEs at a range of heights spanning over the lower solar atmosphere. In addition, it was also found that such localized magnetic structures can support the propagation of both compressive (see, for example, Bloomfield et al. 2004; Jess et al. 2012) and incompressible (Morton et al. 2014) waves, for example kink and Alfvén waves (Erdélyi & Fedun 2007). Indeed, McAteer et al. (2002) have shown the presence of long-period waves in chromospheric bright points that are not consistent with the observational signatures expected for acoustic waves but rather for MHD waves. Among the many types of MHD modes, kink waves have been observed at different regions in the solar atmosphere; from the lower photosphere (e.g., Keys et al. 2011) to the chromosphere (e.g., Jafarzadeh et al. 2013) and the corona

(e.g., Tomczyk et al. 2007; McIntosh et al. 2011). It is generally believed that kink waves are continuously generated thanks to the photospheric granular buffeting action. In this regard, an observational proof supporting this scenario was recently provided by Stangalini et al. (2014), who observed the presence of several subharmonics in the kink-like oscillations of SSMEs in the photosphere, with a fundamental period that is consistent with that of the photospheric granular timescale. Furthermore, the reported presence of subharmonics can be regarded as the signature of a chaotic excitation (Sander & Yorke 2009, 2010). More recently, Stangalini et al. (2015, hereafter Paper I), using high-resolution simultaneous observations at different heights of the solar atmosphere, observed the propagation of kink waves from the photosphere to the chromosphere.

In this work, we advance in this field by studying the temporal orientation of the velocity vector of kink perturbations. We consider the same 35 magnetic elements that were analyzed in Paper I, and investigate their horizontal motion (see Figure 1). We assume that the studied features are chromospheric SSMEs. This assumption is based on the presence of circular polarization signals at their base in the photosphere, and the large coherence between the oscillatory signals in the photosphere and chromosphere reported in Paper I.

2. Data Set and Methods

The data set employed in this work was acquired on 2011 August 6 with the Swedish Solar Telescope (Scharmer et al. 2003). The obtained data consists of a series of chromospheric broadband images centered at the core of the Ca II H line at 396.9 nm on a quiet Sun region at disk center. The estimated formation height of the spectral line is 700 km above the photosphere (Jafarzadeh et al. 2017). In Paper I, the SSMEs were tracked to study kink wave propagation from the

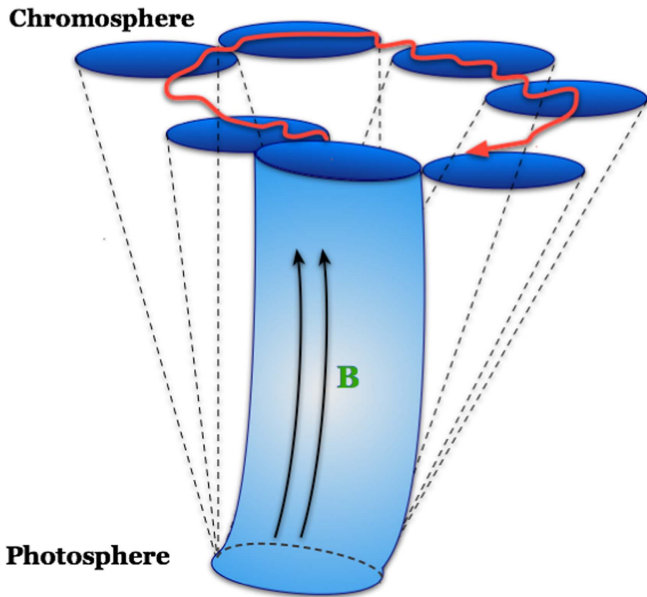


Figure 1. Cartoon of the typical displacement of a SSME as measured in the solar chromosphere. A low-frequency helical displacement is superimposed on a high-frequency kink-like oscillation (red line).

photosphere to chromosphere. The set of magnetic elements constituted a collection of the longest-lived ones for which a chromospheric counterpart was confirmed by the visual inspection of Ca II H data. Therefore, this work may be regarded as a continuation of Paper I, but with a focus only on the chromosphere, where the magnetic elements are not forced directly to move around by the solar photospheric granulation and, perhaps even more importantly, they are free to oscillate.

The observation started at 07:57:39 UT and lasted for 47 minutes with a cadence of the spectral scans of 28 s (100 spectral scans). The pixel scale was 0.034 arcsec/pixel for the Ca II H filter. The spatial resolution of the images is equivalent to ~ 120 km in the solar photosphere. The standard calibration procedure, including the multi-object multi-frame blind deconvolution (van Noort et al. 2005) restoration aimed at limiting seeing-induced distortions, can be seen in the following images.

In Figure 2, we show a typical example of images in the core of the Ca II H line (upper panel), where the selected magnetic elements are highlighted by yellow circular symbols. In the same figure (lower panel), we also show a photospheric image of the same field of view taken at 395.3 nm (1 Å bandpass).

The identification of the chromospheric features was performed on each Ca II line core image of the available series by applying a procedure based on determining the center of mass of the intensity distribution in windows of area 10×10 pixel² encompassing the intensity enhancement co-spatial to the photospheric feature. For more detailed information on the identification, we refer the reader to Paper I. At each time step, the two components of the horizontal velocity of each identified magnetic element were determined as the time derivatives of the measured position of the feature.

2.1. Empirical Mode Decomposition (EMD)

In order to isolate the low-frequency evolution of the horizontal velocity vector and to filter out the high-frequency noise that can affect the signals, we applied the method of

EMD. EMD was introduced by Huang et al. (1998) as a technique for the regularization of a signal before the application of Hilbert transform. EMD decomposes a signal into a set of intrinsic mode functions (IMFs), which represent different oscillations at a local level. For a complete introduction to EMD and its application to solar physics data, we refer the reader to Terradas et al. (2004). Another example of a solar application involving MHD waves is by Morton et al. (2011).

In contrast to the Fourier method, that applies to rigorously stationary and linear time series, the characteristic timescales of the IMFs preserve all the non-stationarities of the signal as well as its nonlinearities. Indeed, an IMF is defined as a local mode, which satisfies the following conditions: (i) the number of extrema and zero-crossing should be equal or differ at the most by one, and (ii) at any time the mean value of the upper and lower envelopes, defined by the local maxima and minima, respectively, is zero. It is important to point out that the EMD technique does not require any a priori assumption, as the decomposition is being based on the data itself. As a result of applying the EMD technique, the original signal $v(t)$ is decomposed into a set of IMFs and a residue R , so that one can write:

$$v(t) = \sum_{i=1}^n IMF_i(t) + R(t). \quad (1)$$

It is important to underline that the EMD method, based on the extraction of the energy contained in the intrinsic timescales of the signal, can be applied successfully to non-stationary signals. These properties make the method very attractive for solar MHD wave research.

Note that, in contrast to Fourier transforms, EMD does not require any transformation of the signal, thus preserving the original nonlinearities (if any) of the process.

The EMD technique has already been employed to study kink waves in small-scale magnetic elements and more information about the application of this technique can be found in Stangalini et al. (2014).

Distinct from earlier works available in the literature, in this study we examine the slow temporal changes of the orientation of the velocity vector of the magnetic elements, instead of the oscillations of its magnitude. To do this, we apply EMD to both components of the horizontal velocity of each magnetic element investigated, and extract the low-frequency part of the signal as in Figure 3. Each velocity signal is decomposed into five IMFs. The first IMF containing the high-frequency part of the signal is then neglected, while the following IMFs are used to reconstruct the signal itself. Indeed, the first IMF mostly captures noise (Flandrin et al. 2004) and high-frequency perturbations due for instance to intergranular turbulence (Jafarzadeh et al. 2014).

3. Results

3.1. A Case Study

Before showing the results derived from the analysis of the 35 magnetic elements selected, in this section we analyze in detail a case study in order to highlight the key points. This is also done to better describe the methods used, and provide a more detailed insight into the results for a particular yet representative case. In this regard, we have selected the magnetic element represented by a red star in Figure 2. This

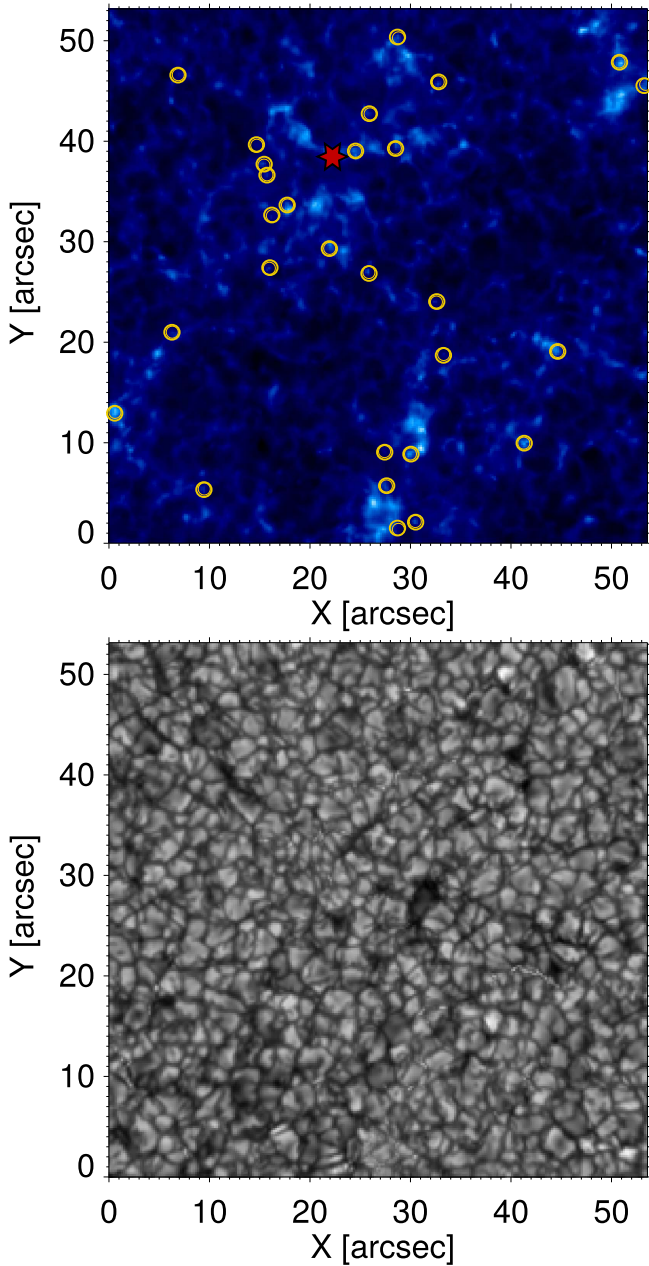


Figure 2. Example of the data analyzed in this study. Top: chromospheric Ca II H broadband image. The 35 longest-lived SSMEs used in the analysis are highlighted by yellow circles. The red star (upper panel) identifies the case study analyzed in the text. Bottom: image of the same FoV taken at 395.3 nm with 1 Å bandpass.

SSME is tracked with an automated procedure that tracks the center-of-mass of the intensity distribution of the element itself. After the automated procedure the tracking is verified by visual inspection, as for all other SSMEs studied in this work. The trajectory $s(t)$ of the selected SSME in time is shown in Figure 4. Superimposed on the trajectory of the SSME, we plot the velocity vector obtained from the derivative of $s(t)$. The horizontal velocity is characterized by a broad spectrum of oscillations (see the right panel in Figure 4), from ~ 1 –2 mHz up to ~ 10 mHz. This is the case for both components of the horizontal velocity vector, which display a good agreement. In the same panel we also overplot the coherence (smoothed with

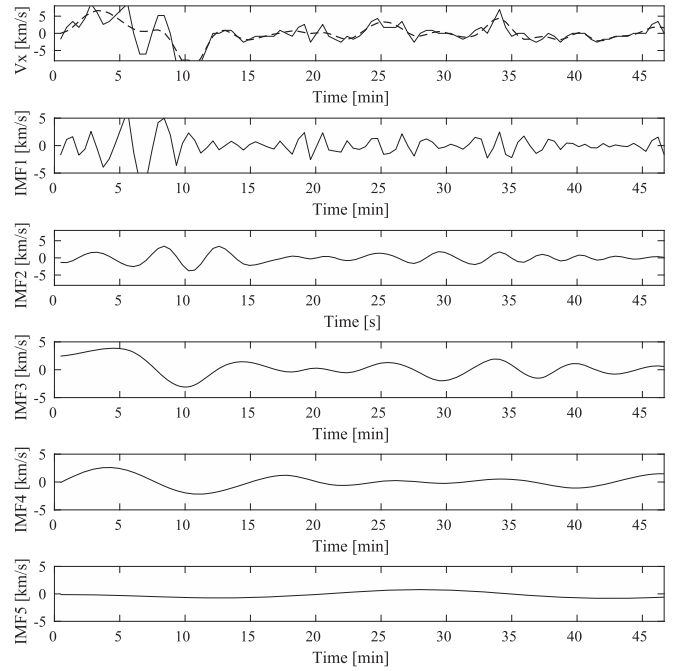


Figure 3. Example of EMD of one of the components of the horizontal velocity vector of a SSME tracked in this study. The black solid line in the upper panel represents the original time series, while the dashed line represents the filtered signal. The subsequent panels show the IMFs that decompose the signal.

an averaging window 3 points wide) between the two spectra defined by

$$C(\nu) = \frac{|P_{xy}(\nu)|^2}{P_{xx}(\nu)P_{yy}(\nu)}, \quad (2)$$

where $P_{xy}(\nu)$ is the cross spectral density between v_x and v_y (the two components of the horizontal velocity vector $\dot{s}(t)$), and $P_{xx}(\nu)$ and $P_{yy}(\nu)$ are the power spectral densities of v_x and v_y , respectively.

It is interesting to note that the visual inspection of the trajectory itself already gives the impression of a rotation of the displacement vector. This appears more evident in Figure 5, where we plot the EMD filtered horizontal velocity vector $\dot{s}(t)$ in polar coordinates. The color scale in this (and any further) polar plot encodes the temporal evolution (from dark red to light yellow). These plots show that the orientation of the velocity vector $\dot{s}(t)$ changes smoothly in time and does not present jumps in the orientation itself. In other words, there exists a long-term memory of the process, which determines the evolution of the velocity vector. This can be also seen in the right panel of the same figure, where we plot the same information contained in the polar plot previously described, but we have unrolled it along the time axis. This graph clearly displays a rotation of the velocity vector in time. The rotation of $\dot{s}(t)$ takes place over the first 30 minutes only, while in the remaining fraction of the lifetime of the SSME, no rotation of the velocity vector, is observed. Referring to the trajectory of the SSME shown in the upper right panel of Figure 4, we see that the first part of the elements lifetime is marked by a distinguishable helical motion of the flux tube (trajectory points lying in the right half part of the plot). In this regard, the polar maps of Figure 5 offer a much clearer visualization of this behavior, thus in the rest of this work we will focus on such polar plots of the velocity vector. This process is schematically depicted in Figure 1.

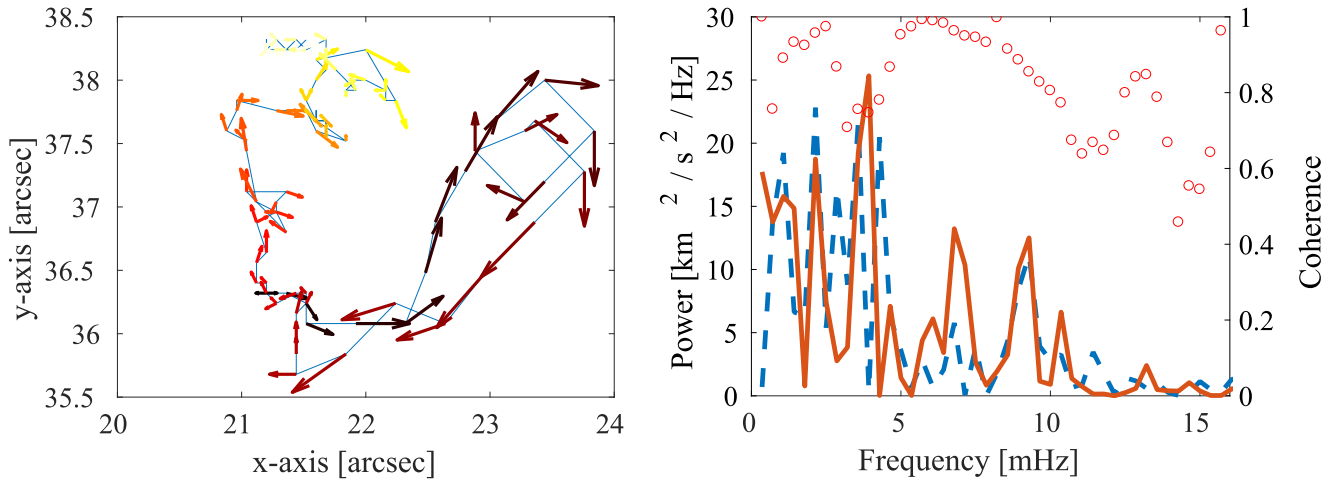


Figure 4. Left panel: trajectory in time $s(t)$ of the magnetic element marked with a red star in Figure 2, with the velocity vector $\dot{s}(t)$ superimposed. The colors encode the temporal evolution, from dark red to light yellow. Right panel: periodogram of the two components of $\dot{s}(t)$ (v_x : red line, v_y : dashed blue line). In the same plot, we also show with the coherence spectrum between these two components, smoothed with an averaging window 3 points wide (red circles).

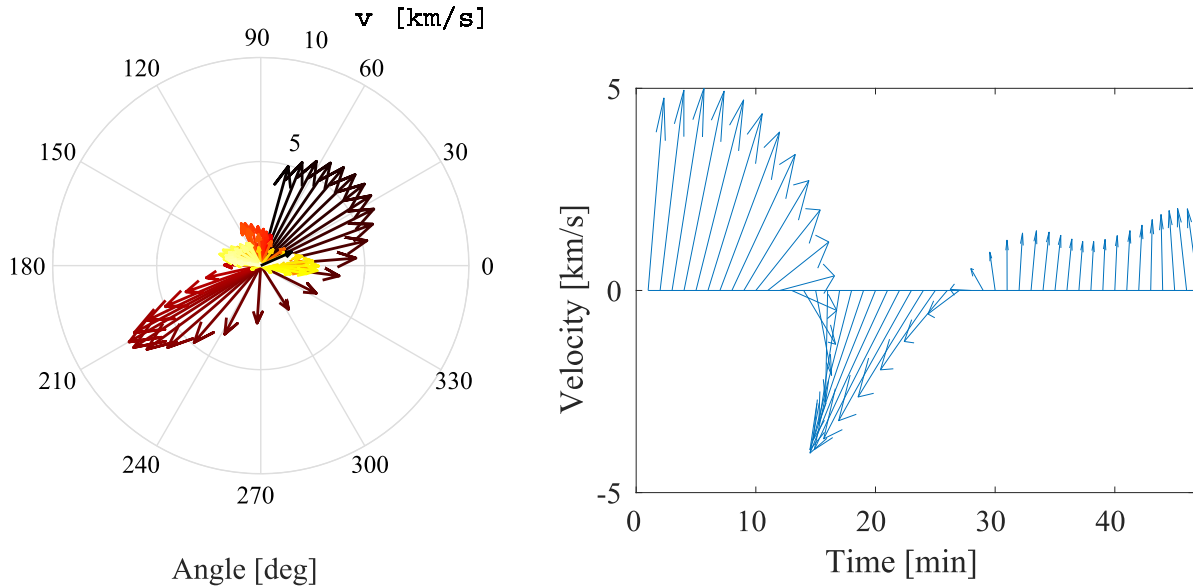


Figure 5. Left panel: evolution of the velocity vector $\dot{s}(t)$ of the element marked as a star in Figure 2 as a function of time in polar coordinates. The color scale encodes the temporal evolution from dark red to yellow. The latitudinal circles represent the magnitude of the vector in km s⁻¹. Right panel: velocity vector (orientation and magnitude) as a function of time.

The results in Figures 4–6 show that the rotation of the horizontal velocity vector is a temporally coherent process that occupies a significant fraction of the lifetime of the selected magnetic element (~ 30 minutes). Indeed, the orientation of the velocity vector is not randomly distributed in space, but follows a helical evolution that suggests a phase lag between v_x and v_y .

3.2. Analysis of the SSMEs Sample

In Figure 6, we visualize the evolution of the horizontal velocity vector in the polar coordinates (after EMD decomposition) of a sample of 6 magnetic elements selected among the 35 elements. The compass plots reveal that the horizontal velocity oscillations of the magnetic elements are not randomly oriented in space, but follow nearly helical trajectories (i.e., the velocity vector evolves smoothly in time, without sudden changes of its orientation). In other words, the horizontal velocity of the magnetic elements appears elliptically polarized, for a significant fraction of the SSMEs lifetime. In addition to

this, the helical motion of the velocity vector is sometimes seen to revert the direction of its angular motion from clockwise to counterclockwise and vice versa.

In order to give an independent and more quantitative characterization of the elliptical motion of the velocity vector seen in the examples of Figure 6, we estimate the helicity $H(\omega)$ of the velocity perturbations. The helicity can be written as follows (Carbone & Bruno 1997):

$$H(\omega) = \frac{2 \operatorname{Im}[a_x^* a_y^*]}{\omega}, \quad (3)$$

where ω is the frequency, and a_x and a_y are the FFT transforms of v_x and v_y , respectively.

Let us now normalize the helicity as follows:

$$\sigma(\omega) = \frac{\omega H(\omega)}{|a_x|^2 + |a_y|^2}. \quad (4)$$

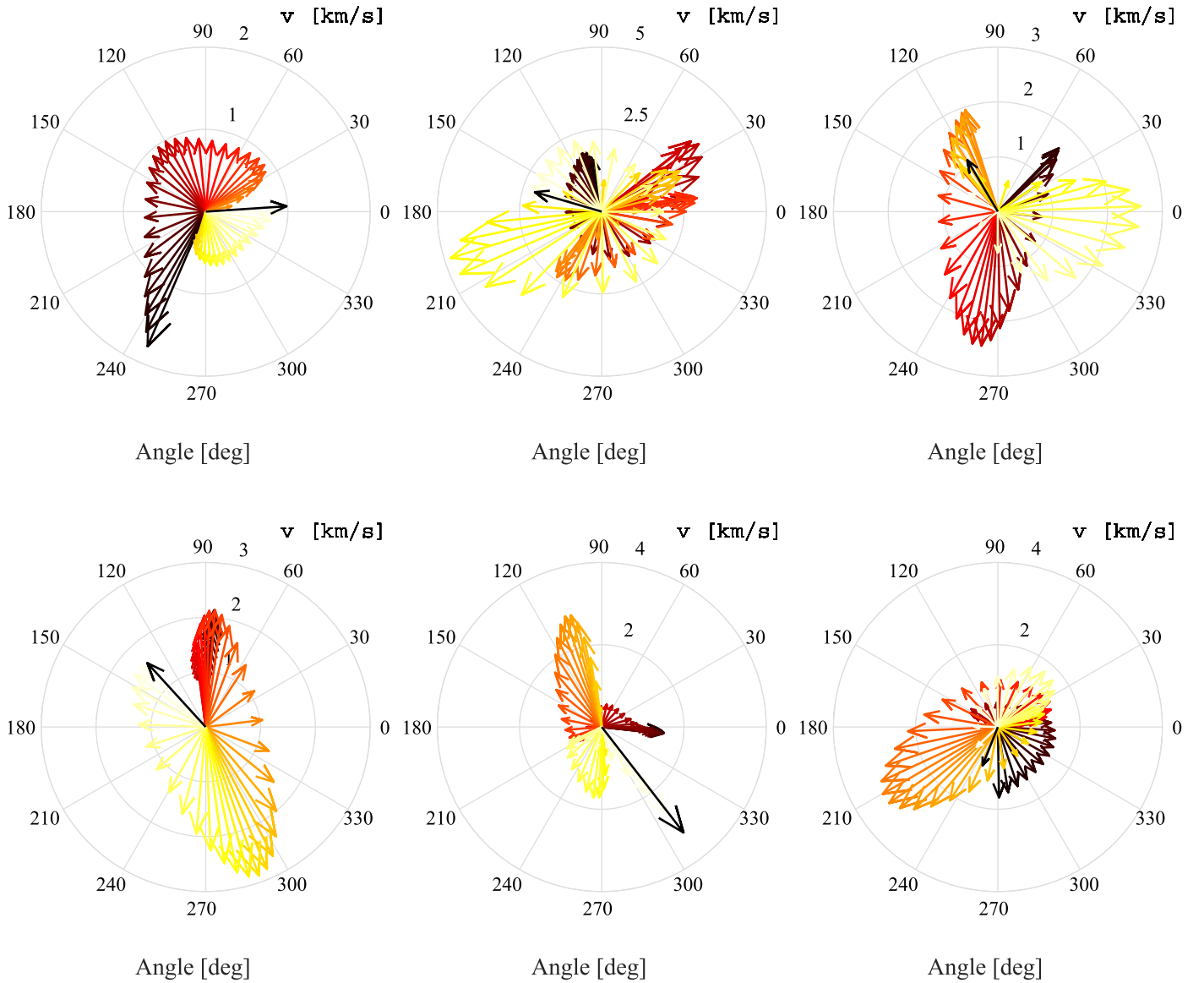


Figure 6. Compass polar diagrams of the horizontal velocity vector of six magnetic elements chosen among the 35 analyzed in this work. The color scale represents the temporal evolution (from dark red to light yellow). The azimuthal lines (latitudes) indicate the magnitude of the horizontal velocity vector in km s^{-1} .

It is worth noting that although the helicity was initially used to study fluctuations in the solar wind, it is clear that Equation (3) can be applied to any time series, as it represents a relation between modes of the Fourier space. The normalized helicity of velocity oscillations of all the 35 magnetic elements is shown in Figure 7 (upper panel). The figure shows that a large number of modes in the spectrum display a $\sigma > 0$ or $\sigma < 0$. This means, clockwise or counterclockwise rotation, respectively, and confirms the earlier derived polarization of the horizontal velocity vector of the magnetic elements in the solar chromosphere. In the same plot, we only show those points with coherence larger than 0.8 and with a cross-power spectral density larger than $300 \text{ km}^4 \text{ s}^{-4} \text{ Hz}^{-1}$ (see bottom panel of the same figure). This threshold is chosen in order to isolate the most prominent peak in the cross-correlation plot in the range of 0–10 mHz. The large coherence also represents a high confidence level, ensuring the reliability of the measurements with a value of 0.8 being a very stringent confidence threshold.

3.3. A Toy Model

With the aim of validating our results, we applied our method to two control simulations: a random process and an oscillatory process, in which two pulses with different directions are superimposed.

In the first case, a random velocity with normal distribution is simulated, see upper panel of Figure 8. In order to apply exactly the same method used with the real data, we filtered the simulated velocity signal employing the EMD (middle panel of Figure 8). The polarization of this filtered signal is then visualized through a compass diagram just like what was done in the case of the observed data. Besides, we simulated a simple oscillatory signal of the form:

$$v_{x,y}(t) = A \sin \frac{2\pi t}{T_1} \sin \frac{2\pi t}{T_2}, \quad (5)$$

where A is the amplitude of the velocity signal, T_1 is period, and T_2 is long-term modulation to simulate a wave train (see lower panel of Figure 8).

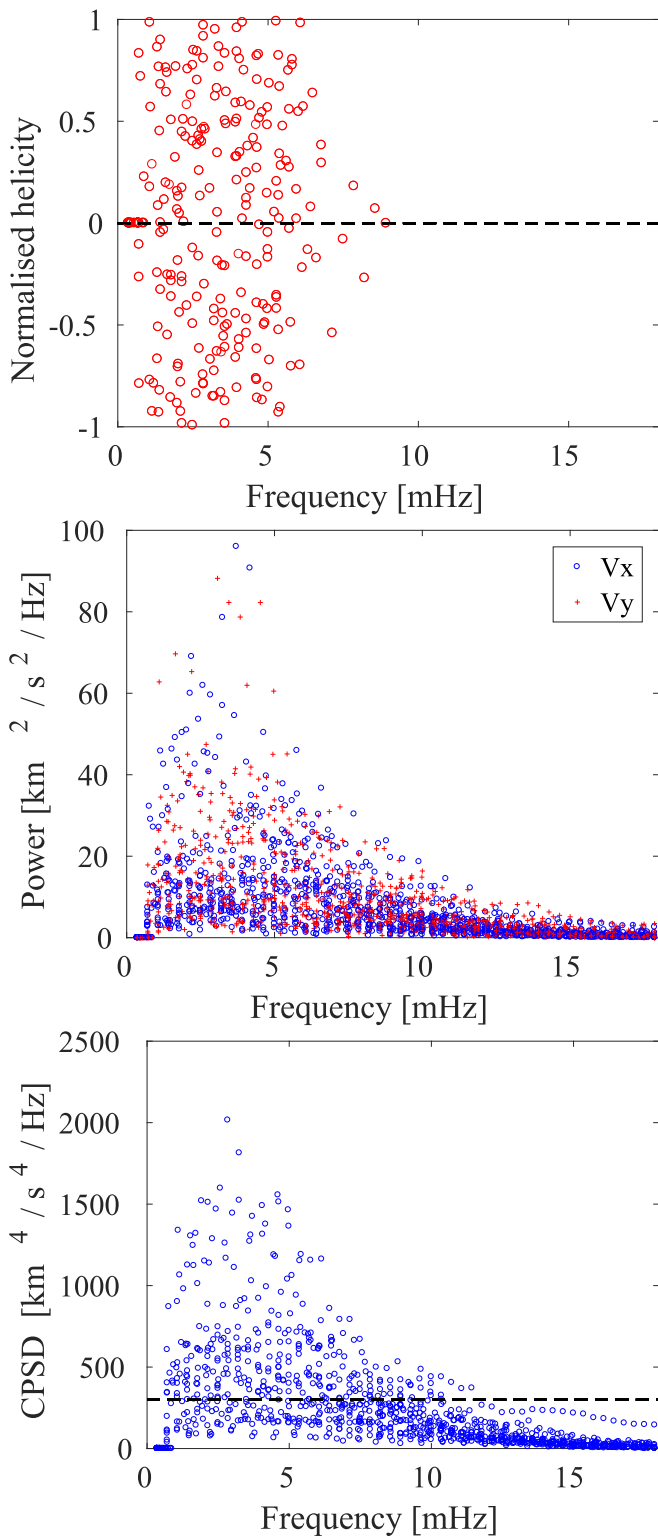


Figure 7. Top: helicity spectrum of the horizontal velocity vector of the 35 magnetic elements investigated. Only Fourier modes with a coherence larger than 0.8 and a cross-correlation larger than the threshold are indicated in the bottom panel of this figure. Middle: power spectra of the two components of the horizontal velocity for all the elements selected. Only modes with a coherence larger than 0.8 are shown. Bottom: cross-power spectral density computed between the two components of the velocity vector. Only modes with a coherence larger than 0.8 are shown. The horizontal dashed line represents the threshold used in the top panel of this figure.

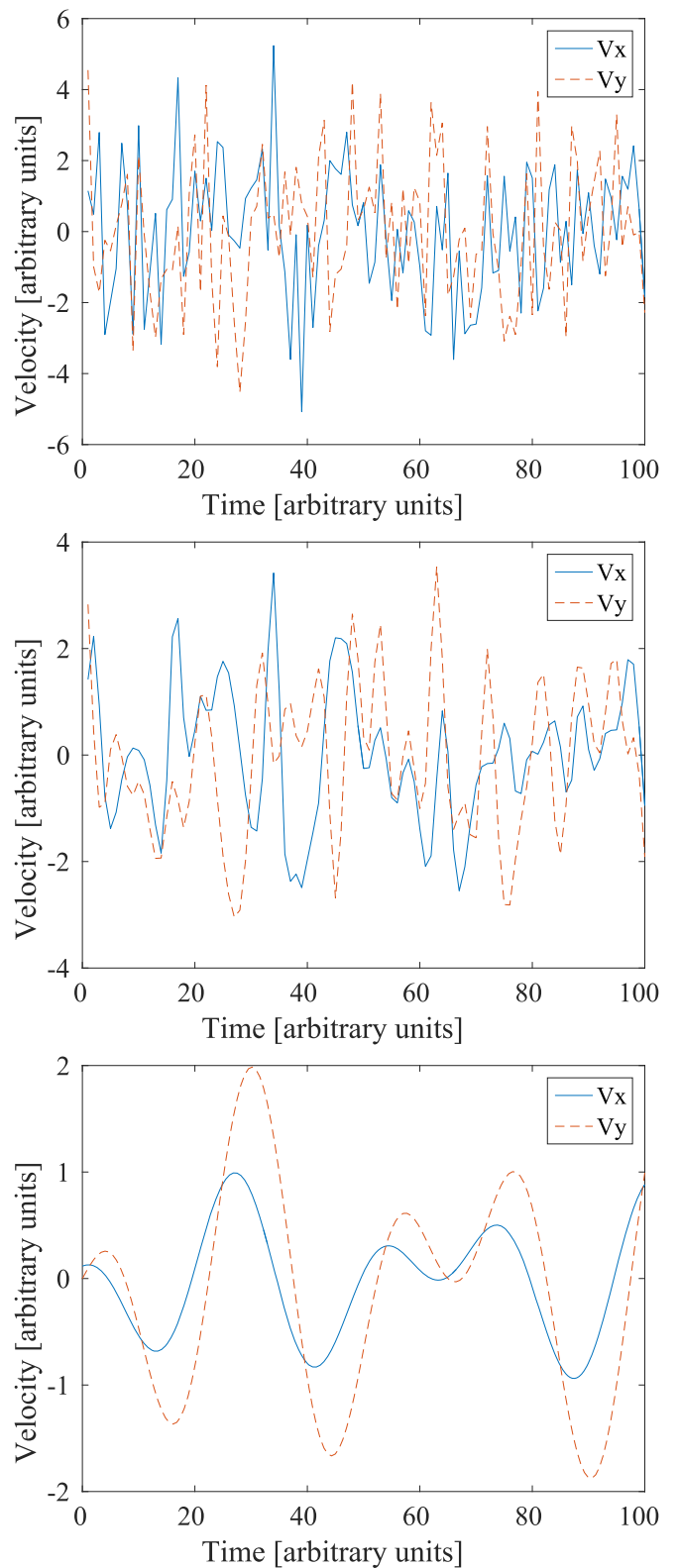


Figure 8. Top panel: simulated random velocity signal with normal distribution. Middle panel: same velocity signal filtered using the EMD. Bottom panel: simulated velocity signal made with two superimposed pulses in the x and y direction respectively. The two pulses are exactly the same, but slightly shifted and with different amplitudes.

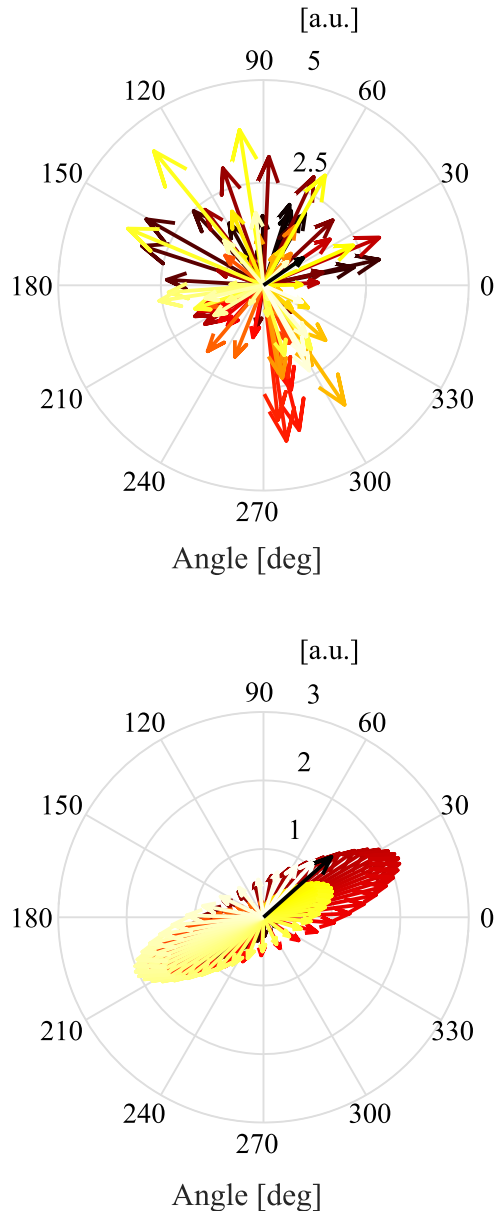


Figure 9. Compass diagram of the simulated random velocity signal with normal distribution (top), and the one obtained from a superposition of two velocity pulses in two perpendicular directions (bottom).

The compass diagrams of the two simulated velocity signals are shown in Figure 9. In comparison with the oscillatory signal, the random process does not display polarization (top panel). This confirms the validity of our findings. In contrast, the same kind of polarization of the compass diagram is indeed obtained with a superposition of two pulses in different directions (bottom panel).

4. Discussion and Conclusion

The results presented above show low-frequency ($<5\text{--}6$ mHz) kink-like horizontal oscillations of SSMEs in chromosphere with an elliptic polarization of the velocity vector (see Figure 1 for a schematic representation) observed in all of the 35 SSMEs analyzed in the quiet chromosphere. This frequency range is inconsistent with residual seeing aberrations. In this regard, it is worth noting that the orientation of the velocity

vector in time is not random, but reflects a long-term coherency (several minutes). This timescale is by far longer than the typical timescale of atmospheric turbulence (a few ms). Indeed, Stangalini et al. (2016) have deeply analyzed the decorrelation timescales of adaptive optic residual aberrations using on-sky data, and found them to be of the order of 10–20 ms at visible wavelengths (~ 630 nm). This is consistent with independent works in the literature (see, for example, Davis & Tango 1996). Since the turbulence timescales at the Ca II H wavelengths is even shorter (Roddier 1999), we expect a large reduction of the turbulence coherence time, resulting in a turbulence timescale that is even more inconsistent with that of the process observed here. In addition, since the seeing residual aberrations are a random process, one cannot expect any coherent helical motion of the velocity vector from them.

We would like to comment here on the selection of the sample of SSMEs used in this work. Indeed, the choice of the 35 chromospheric magnetic elements analyzed in Paper I, and in this work, is twofold. First, in the chromosphere the magnetic elements are not forced by the surrounding granular flows, thus the horizontal oscillations measured represent their free oscillation and cannot be misinterpreted as the result of the photospheric forcing itself. Second, in Paper I, it was shown that these elements, observed as brightenings in Ca II H, could be unambiguously associated with magnetic elements. Indeed, here it was possible to clearly locate their corresponding circular polarization signals in the photosphere. This fact, in particular, ensures that the brightenings tracked in the Ca II H data represents that of the chromospheric counterpart of photospheric magnetic features.

A helical motion was already observed in solar spicules by, e.g., De Pontieu et al. (2007), Suematsu et al. (2008). For a complete historical review of the observations of waves in solar spicules we refer the reader to Zaqarashvili & Erdélyi (2009). In the attempt to put this observational aspect in a theoretical framework, Zaqarashvili & Skhirtladze (2008) have shown that the superposition of random photospheric pulses with different orientations may easily explain the observed polarized motion. Indeed, these authors demonstrated that each photospheric forcing pulse can excite a kink wave in the flux tube whose polarization plane depends on the pulse itself. The superposition of two or more pulses, and then of different kink waves polarized in different planes, may give rise to a complex oscillation of the flux tube and set up helical waves. Our observational results are perfectly in agreement with this theoretical prediction, and can be seen as the counterpart of the polarized motion in spicules observed by De Pontieu et al. (2007) and those predicted by Zaqarashvili & Skhirtladze (2008). Moreover, our results confirm that the superposition of different driving pulses, with different amplitudes, is effective in generating kink waves with elliptical polarization that can propagate from the photosphere to the chromosphere, where they still maintain their kink-like oscillatory behavior.

In this regard, Zaqarashvili & Erdélyi (2009) noted that the flux tube expansion in the solar chromosphere may hamper the propagation of kink waves. However, our results show that, as soon as the heights spanned by the Ca II H are concerned (~ 700 km above the photosphere), the kink wave appears as the most plausible explanation for the observed oscillations (Erdélyi & Fedun 2007).

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References

- Berger, T. E., Löfdahl, M. G., Shine, R. S., & Title, A. M. 1998, *ApJ*, 495, 973
 Bloomfield, D. S., McAteer, R. T. J., Lites, B. W., et al. 2004, *ApJ*, 617, 623
 Bonet, J. A., Cabello, I., & Sánchez Almeida, J. 2012, *A&A*, 539, A6
 Carbone, V., & Bruno, R. 1997, *ApJ*, 488, 482
 Chitta, L. P., van Ballegoijen, A. A., Rouppe van der Voort, L., DeLuca, E. E., & Kariyappa, R. 2012, *ApJ*, 752, 48
 Davis, J., & Tango, W. 1996, *PASP*, 108, 456
 De Pontieu, B., Erdélyi, R., & James, S. P. 2004, *Natur*, 430, 536
 De Pontieu, B., McIntosh, S. W., Carlsson, M., et al. 2007, *Sci*, 318, 1574
 Edwin, P., & Roberts, B. 1983, *SoPh*, 88, 179
 Erdélyi, R., & Fedun, V. 2007, *Sci*, 318, 1572
 Fedun, V., Shelyag, S., & Erdélyi, R. 2011, *ApJ*, 727, 17
 Flandrin, P., Gonçalves, P., & Rilling, G. 2004, *ISPL*, 11.2, 112
 Giannattasio, F., Berrilli, F., Biferale, L., et al. 2014, *A&A*, 569, A121
 Giannattasio, F., Del Moro, D., Berrilli, F., et al. 2013, *ApJL*, 770, L36
 Hasan, S. S., Kalkofen, W., van Ballegoijen, A. A., & Ulmschneider, P. 2003, *ApJ*, 585, 1138
 Huang, N. E., Shen, Z., Long, S. R., et al. 1998, *RSPSA*, 454, 903
 Jafarzadeh, S., Cameron, R. H., Solanki, S. K., et al. 2014, *A&A*, 563, A101
 Jafarzadeh, S., Rutten, R. J., Solanki, S. K., et al. 2017, *ApJS*, 229, 11
 Jafarzadeh, S., Solanki, S. K., Feller, A., et al. 2013, *A&A*, 549, A116
 Jefferies, S. M., McIntosh, S. W., Armstrong, J. D., et al. 2006, *ApJL*, 648, L151
 Jess, D. B., Shelyag, S., Mathioudakis, M., et al. 2012, *ApJ*, 746, 183
 Keys, P. H., Mathioudakis, M., Jess, D. B., et al. 2011, *ApJL*, 740, L40
 Lagg, A., Solanki, S. K., Riethmüller, T. L., et al. 2010, *ApJL*, 723, L164
 McAteer, R. T. J., Gallagher, P. T., Williams, D. R., et al. 2002, *ApJL*, 567, L165
 McIntosh, S. W., de Pontieu, B., Carlsson, M., et al. 2011, *Natur*, 475, 477
 Morton, R. J., Erdélyi, R., Jess, D. B., & Mathioudakis, M. 2011, *ApJL*, 729, L18
 Morton, R. J., Verth, G., Hillier, A., & Erdélyi, R. 2014, *ApJ*, 784, 29
 Morton, R. J., Verth, G., Jess, D. B., et al. 2012, *NatCo*, 3, 1315
 Musielak, Z. E., Rosner, R., & Ulmschneider, P. 1989, *ApJ*, 337, 470
 Musielak, Z. E., & Ulmschneider, P. 2003, *A&A*, 400, 1057
 Nutto, C., Steiner, O., Schaffnerberger, W., & Roth, M. 2012, *A&A*, 538, A79
 Roberts, B. 1983, *SoPh*, 87, 77
 Roberts, B., & Webb, A. R. 1978, *SoPh*, 56, 5
 Roddier, F. 1999, *Adaptive Optics in Astronomy* (Cambridge: Cambridge Univ. Press)
 Sander, E., & Yorke, J. A. 2009, arXiv:0910.3570
 Sander, E., & Yorke, J. A. 2010, arXiv:1002.3363
 Scharmer, G. B., Dettori, P. M., Löfdahl, M. G., & Shand, M. 2003, *Proc. SPIE*, 4853, 370
 Spruit, H. C. 1981, *A&A*, 98, 155
 Stangalini, M. 2014, *A&A*, 561, L6
 Stangalini, M., Consolini, G., Berrilli, F., De Michelis, P., & Tozzi, R. 2014, *A&A*, 569, A102
 Stangalini, M., Giannattasio, F., & Jafarzadeh, S. 2015, *A&A*, 577, A17
 Stangalini, M., Pedichini, F., Ambrosino, F., Centrone, M., & Del Moro, D. 2016, *Proc. SPIE*, 9909, 99097P
 Steiner, O., Grossmann-Doerth, U., Knoelker, M., & Schüssler, M. 1998, *ApJ*, 495, 468
 Suematsu, Y., Ichimoto, K., Katsukawa, Y., et al. 2008, in ASP Conf. Ser. 397, First Results From Hinode, ed. S. A. Matthews, J. M. Davis, & L. K. Harra (San Francisco, CA: ASP), 27
 Terradas, J., Oliver, R., & Ballester, J. L. 2004, *ApJ*, 614, 435
 Tomczyk, S., McIntosh, S. W., Keil, S. L., et al. 2007, *Sci*, 317, 1192
 van Noort, M., Rouppe van der Voort, L., & Löfdahl, M. G. 2005, *SoPh*, 228, 191
 Vigeesh, G., Fedun, V., Hasan, S. S., & Erdélyi, R. 2012, *ApJ*, 755, 18
 Zaqarashvili, T., & Erdélyi, R. 2009, *SSRv*, 149, 355
 Zaqarashvili, T. V., & Skhirtladze, N. 2008, *ApJL*, 683, L91